SCUBA-2 OBSERVATIONS OF NEARBY SPIRAL GALAXIES

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Why do we need 850 µm data?

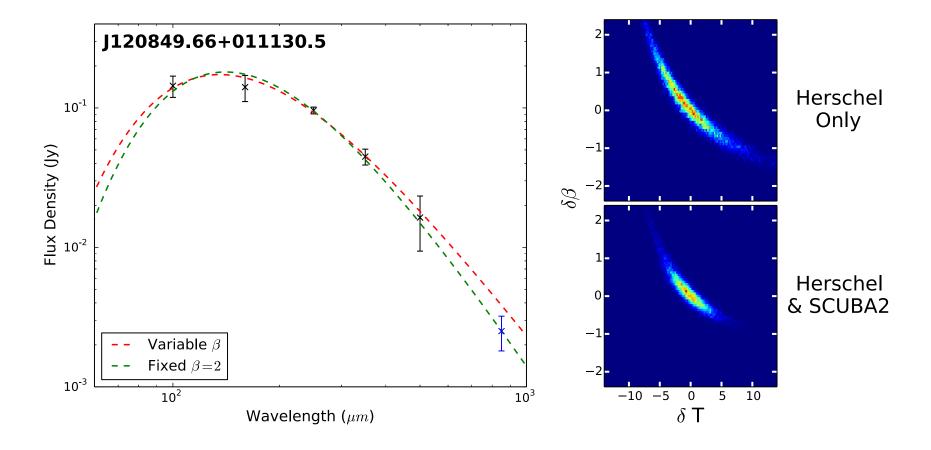
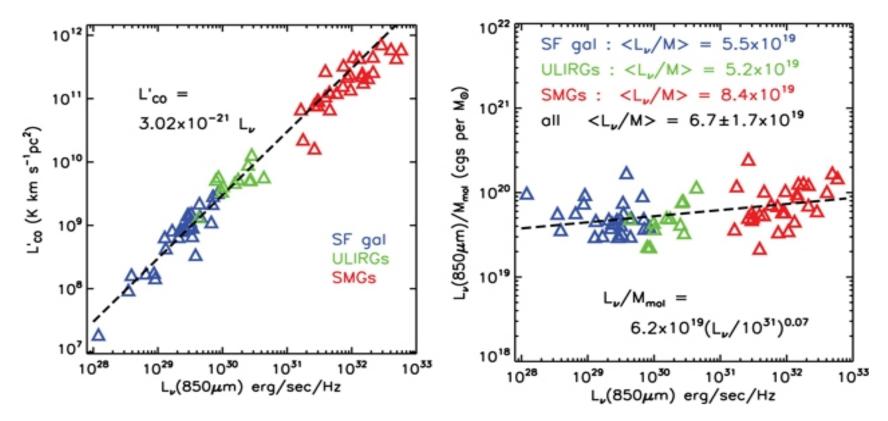
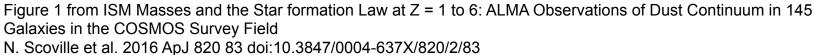


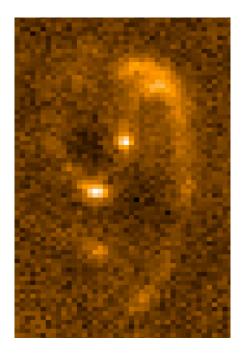
Figure from JINGLE proposal 2015; by Ilse de Looze

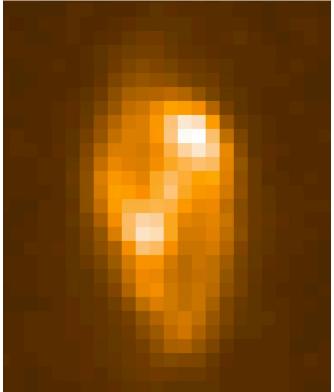
μ m flux as proxy for gas mass





Higher resolution: SCUBA-2 vs Herschel (NGC 3627)





450 micron (cropped) JCMT+SCUBA2 NGLS 500 micron Herschel+SPIRE KINGFISH (Kennicutt et al. 2011, PASP)

The JCMT Nearby Galaxies Legacy Survey: an HI-flux selected sample

- 155 galaxies between 2 and 25 Mpc
- HI flux > 6 Jy km/s
- Sample described in Wilson et al. 2012, MNRAS
 - 47 SINGS galaxies (Kennicutt et al. 2003)
 - 18 HI brightest Irr and E galaxies (HI flux > 3 Jy km/s) + 18 randomly selected spirals in Virgo Cluster
 - random selection of 72 field galaxies
 - D₂₅ < 5'
 - Randomly select 18 galaxies in each of 4 morphology bins (E, early S, late S, Irr)

JCMT data

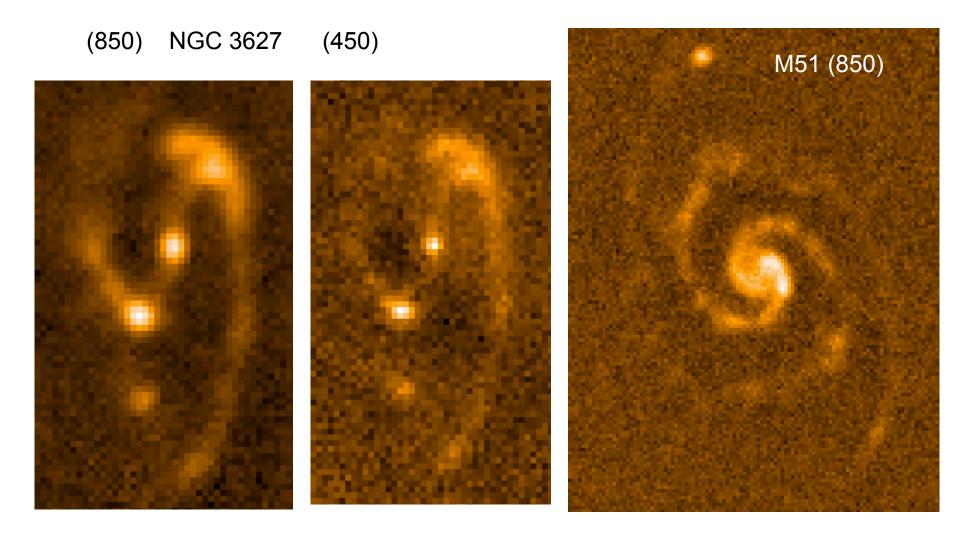
SCUBA-2 850 and 450 micron imaging (100 hours)

- Area covered is roughly $D_{25}/2$
- Sensitivity goal 1.8 mJy/beam at 850 μm in center of map
- **48 spiral galaxies observed** (25 from SINGS sample)
- Only ~40% of SINGS galaxies observed in sufficiently good weather that 450 μm data usable

CO J=3-2 data cubes (256 hours)

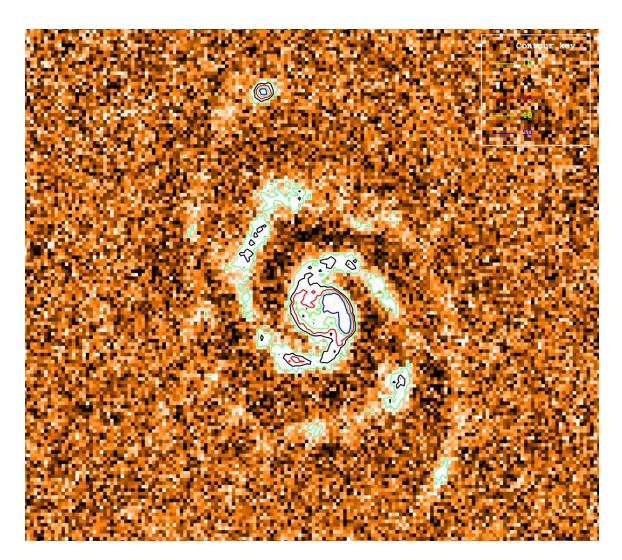
- Area covered is D₂₅/2 for all 155 galaxies
 - Plus follow-up PI program on additional Virgo spirals
- Sensitivity 19 mK at 20 km/s resolution rms
- Equivalent to $Av = 1 \text{ mag or } 2x10^{21} \text{ H/cm}^2 \text{ rms}$
- Detect ~45% of spirals, 20% of ellipticals, almost no irregulars

SCUBA-2 images of nearby galaxies



Wilson, Parkin, Newton, Meng, He et al., in prep

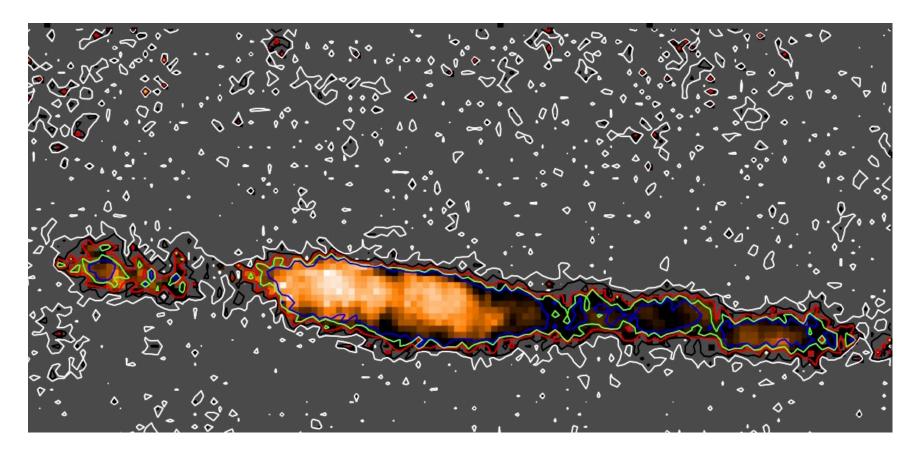
Recent work on SCUBA-2 data



M51 at 850 microns

Contours 10,20,30, 40,50 mJy/beam

CO contamination of 850 micron flux: NGC 4631

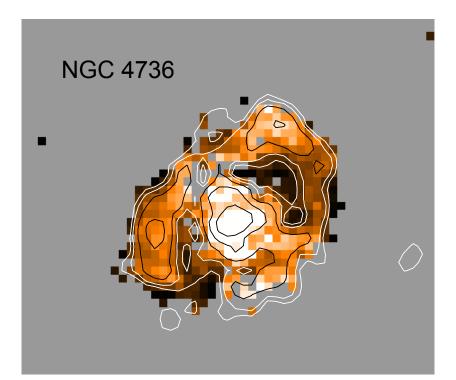


Colour: CO / 850; contours: 850 micron flux (4,8,12,16,20 mJy/beam)

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Significant CO contamination of bright central regions of galaxies

- Average contamination 0.25 for 20 galaxies
 - But significant range:
 - 0.12 in NGC4631
 - 0.35 in M82
- Some indication of higher contamination in starburst or Seyfert nucleus
- Lower contamination in NGC4631might suggest lower contamination in outer portions of disk?

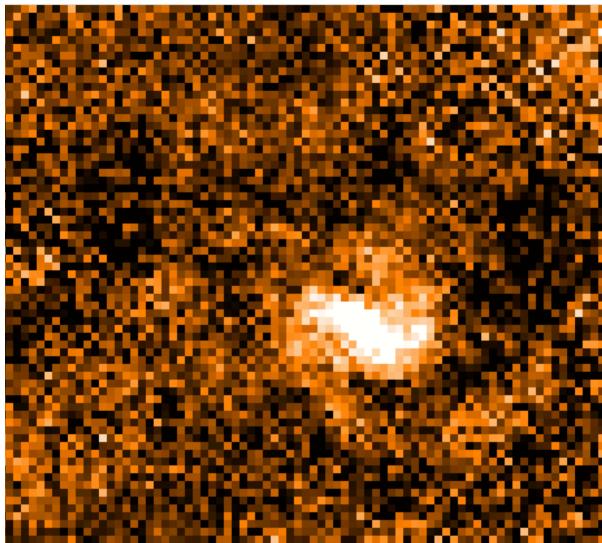


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Color scale % CO contamination, range from 0-50% CO contours 1,2,4,8,16 K km/s

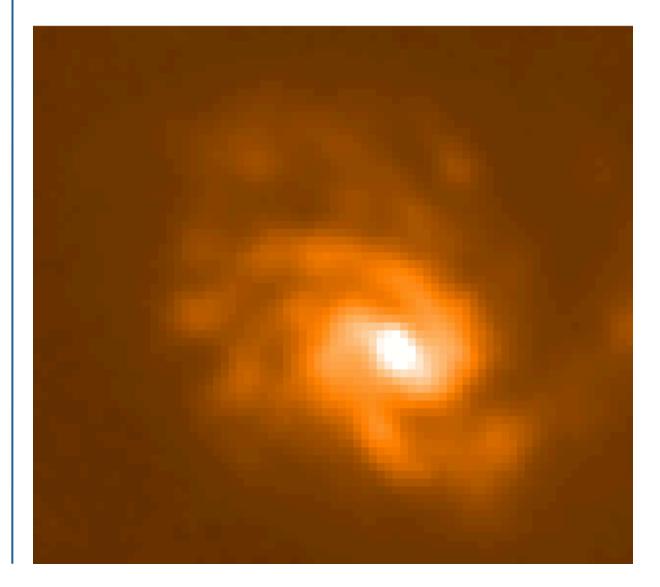
Far-infrared colors: constraining the dust temperature

NGC 4254 850 μm image



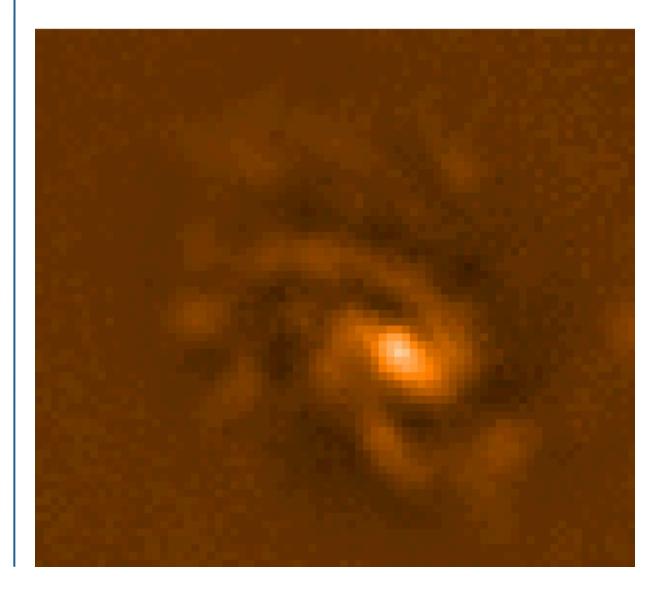
$\begin{array}{c} \text{NGC 4254} \\ \text{250} \ \mu\text{m} \ \text{image} \end{array}$

Original Herschel image from KINGFISH survey, Kennicutt et al. 2010



NGC 4254 250 µm image, skylooped

Herschel image was scaled down so that its brightness was similar to the 850 image; added as a fake source to the SCUBA-2 data time stream; and then reimaged

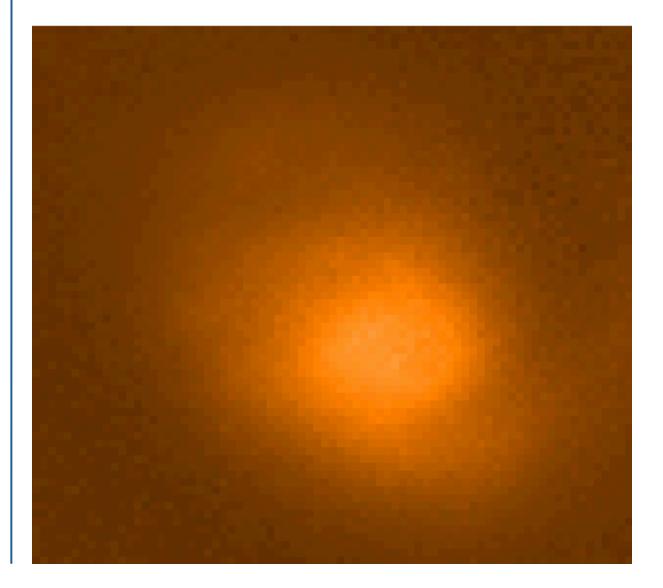


NGC 4254 250 µm flux missed

Flux that is removed in the SCUBA-2 data processing

This image is the difference of the previous two images

At the peak region in the center of the galaxy, at least 1/3 of the flux is gone after skyloop filtering



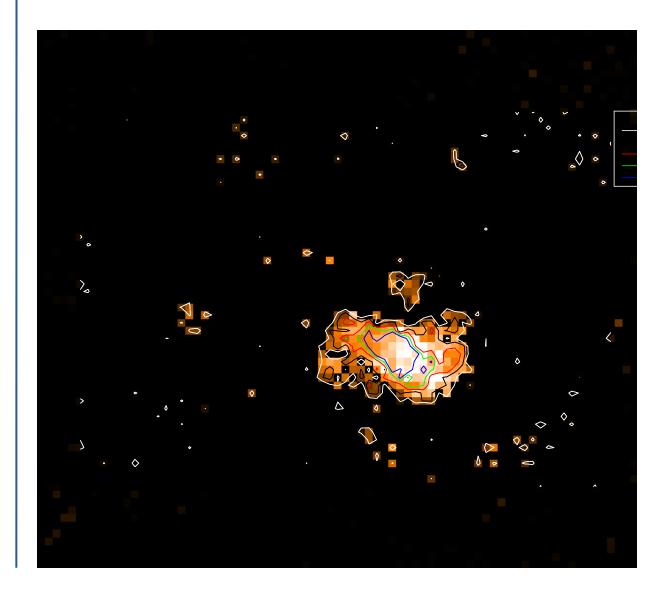
NGC 4254 250/850 ratio

Both images have the same spatial filtering

Color 250/850 flux ratio, scale is 0-40

Ratio map has been clipped at 2sigma in 850 map

Contours are 850 surface brightness, 2,3,4,5,6 sigma contours (1sigma = 0.017 mJy/ square arcsecond)



Conclusions

- Analysis of SCUBA-2 data still at a very early stage
- Significant CO 3-2 contamination in the 850 μm images, especially in galaxy centers
- Comparison with Herschel data should let us constrain dust properties such as temperature

Public link to reduced CO J=3-2 data cubes and maps for complete NGLS: http://www.physics.mcmaster.ca/~wilson/www_xfer/NGLS/

Are you interested in SCUBA-2 data? Contact me!

NGC628	NGC925	NGC2403	NGC2841	NGC2976	NGC3034
NGC3184	NGC3198	NGC3351	NGC3521	NGC3627	NGC3938
NGC4254	NGC4321	NGC4559	NGC4569	NGC4579	NGC4594
NGC4631	NGC4736	NGC4826	NGC5033	NGC5055	NGC5194
IC0750	NGC404	NGC3077	NGC3162	NGC3277	NGC3353
NGC4013	NGC4041	NGC4567	NGC4568	NGC4647	

Note: not all galaxies are detected Red indicates weather too poor for 450 μm observations